

# Tidal Evolution of Close-In Extra-Solar Planets

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**Abstract.** The distribution of eccentricities  $e$  of extra-solar planets with semi-major axes  $a > 0.2$  AU is very uniform, and values for  $e$  are generally large. For  $a < 0.2$  AU, eccentricities are much smaller (most  $e < 0.2$ ), a characteristic widely attributed to damping by tides after the planets formed and the protoplanetary gas disk dissipated. We have integrated the classical coupled tidal evolution equations for  $e$  and  $a$  backward in time over the estimated age of each planet, and confirmed that the distribution of initial  $e$  values of close-in planets matches that of the general population for reasonable tidal dissipation values  $Q$ , with the best fits for stellar and planetary  $Q$  being  $\sim 10^{5.5}$  and  $\sim 10^{6.5}$ , respectively. The current small values of  $a$  were only reached gradually due to tides over the lifetimes of the planets, *i.e.*, the earlier gas disk migration did not bring all planets to their current orbits. As the orbits tidally evolved, there was substantial tidal heating within the planets. The past tidal heating of each planet may have contributed significantly to the thermal budget that governed the planet's physical properties, including its radius, which in many cases may be measured by observing transit events. Here we also compute the plausible heating histories for a few planets with anomalously large measured radii, including HD 209458 b. We show that they may have undergone substantial tidal heating during the past billion years, perhaps enough to explain their large radii. Theoretical models of exoplanet interiors and the corresponding radii should include the role of large and time-variable tidal heating. Our results may have important implications for planet formation models, physical models of “hot Jupiters”, and the success of transit surveys.

**Keywords.** celestial mechanics, planetary systems: formation, protoplanetary disks

## 1. Introduction

As the number of known extra-solar planetary orbits has increased, interesting patterns are emerging. Figure 1 shows the semi-major axes  $a$  and eccentricities  $e$ . Eccentricities of extra-solar planets with  $a > 0.2$  AU average 0.3 and are broadly distributed up to near 1. The distribution of eccentricities is fairly uniform over  $a$ . For example, a Kolmogorov-Smirnov (K-S) test shows that the  $e$  distribution for  $a$  between 0.2 and 1.0 AU matches that for  $a$  between 1.0 and 5.0 AU at the 96% confidence level. However, for close-in extra-solar planets (by which we mean  $a < 0.2$  AU),  $e$  values are smaller, with an average of 0.09, but still large compared to our solar system. For  $a < 0.2$  AU, the K-S test shows agreement at only the 0.1% level compared with planets further out.

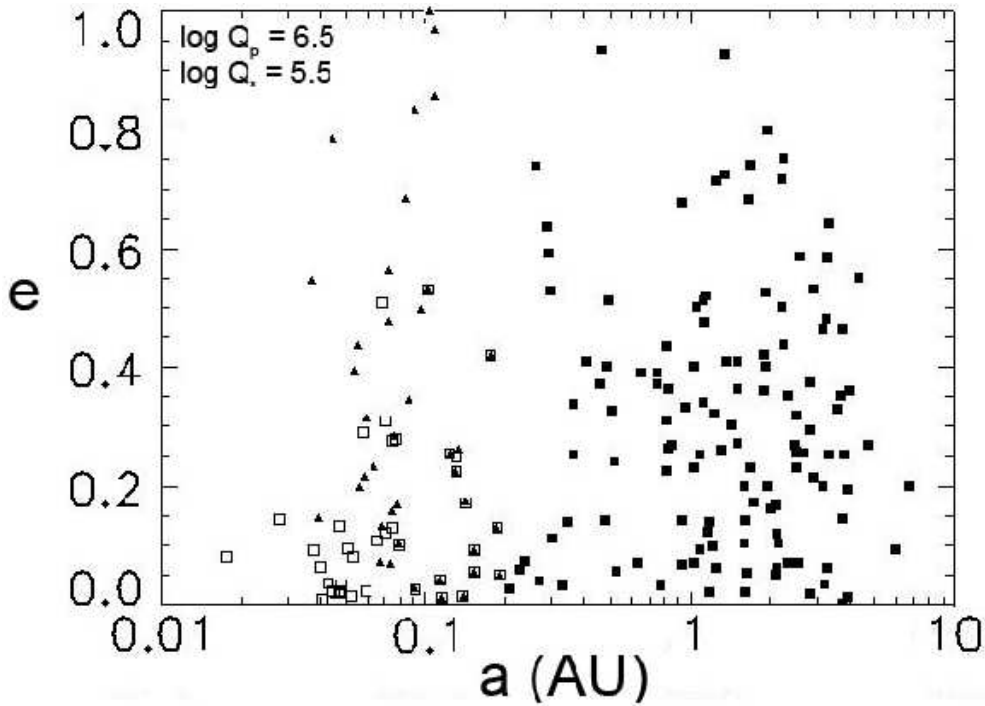
Here we consider the conventional idea that close-in planets began with a distribution of  $e$  similar to that of planets farther out. Because the magnitude of tidal effects falls off very rapidly with increasing  $a$ , Rasio *et al.* (1996) suggested that tides could have reduced  $e$  for close-in planets and not for those farther out.

The implications of this hypothesis for planetary masses  $M_p$ , radii  $R_p$ , and heating rates has been considered in previous studies (Rasio *et al.* (1996); Trilling (2000);

Bodenheimer *et al.* (2003)). Several assumptions have been common. First, the tidal dissipation parameter  $Q_p$  for the planet is usually based on model-dependent estimates based on the tidal evolution of the Jovian satellite system (Yoder & Peale (1981)) or on physical theories of tidal dissipation (Goldreich & Nicholson (1977); Ogilvie & Lin (2004)). Second, effects of tides raised on the star are only partially taken into account. Often the orbit-circularizing effects of tides raised on the host star are ignored, although they can be important. Third, some studies (*e.g.* Ford & Rasio (2006)) assumed that orbital angular momentum is conserved during tidal evolution, which is not true when tides on the star are taken into account. Fourth, previous work ignored the strong coupling of the tidal evolution of  $e$  with that of  $a$ , generally describing the circularization of an orbit in terms of an implicitly exponential “timescale”. In fact, the coupling with  $a$  means that evolution is more complex than a simple exponential damping. Moreover, tidal evolution has resulted in significant inward migration of many close-in planets.

In our study, we avoid these assumptions. We use the full tidal evolution equations to determine possible past orbital change. The equations account for tides raised on the stars by the planet and on the planet by the star, each of which affects changes in both  $e$  and  $a$ . We employ conservative assumptions about  $M_p$  and  $R_p$ , to test the tidal circularization hypothesis. And we consider a wide range of possible  $Q$  values, allowing us to determine which pair of  $Q_p$  and  $Q_*$  values yields the most plausible evolution history.

These calculations also yield the corresponding past tidal heating history for each planet. During the course of the tidal evolution, tidal distortion of the figure of the planet can result in substantial amounts of internal heating at the expense of orbital



**Figure 1.** Distributions of orbital elements. Squares (filled and open) represent the currently observed orbital elements, with the open squares (with  $a < 0.2$  AU) being candidates for significant tidal evolution. Triangles show the initial orbital elements ( $e_{initial}$  and  $a_{initial}$ ) determined by integrating the equations of tidal evolution backward in time to the formation of the planet, using  $Q$  values that gave the best fit of the  $e$  distribution to that of the other planets.

energy, so the heating rate as a function of time is coupled to the evolution of the orbit. In a typical case, tidal heating might have begun modest, but then increased as tides reduced  $a$ . As the tides became stronger, they would circularize the orbit, which in turn would shut down the tidal heating mechanism. The relative strength and timing of these two effects would determine a planet's history, typically with a gradual increase in the heating rate followed by a decrease.

The thermal history of a planet is critical to determining its physical properties. For example, studies of extra-solar planets have considered the effects of heating on their radii, which can be measured directly by transit observations. Heat sources that have been considered in these models include the energy of planetary accretion and radiation from the star, as well as tidal heating (Bodenheimer *et al.* (2003); Mardling (2007)). In many cases the theoretical predictions match the observations reasonably well (Burrows *et al.* (2007)). However, there are notable exceptions. HD 209458 b has been observed by Knutson *et al.* (2007) to have a radius of 1.32 Jupiter radii ( $R_J$ ), which is 10-20% larger than predicted by theoretical modeling Guillot (2005). Similarly, HAT-P-1 b is 10-20% larger than predicted by theory (Bakos *et al.* (2007a)). On the other hand, HAT-P-2 b seems to be smaller than theory would predict (Bakos *et al.* (2007b)), while for GJ 436 b, theory predicts a radius consistent with observation (Gillon *et al.* (2007)).

Theoretical models to date have not taken into account the history of tidal heating for close-in planets, and of course those are the planets most likely to have radii measurable by transits. The tidal heating histories reported here and in Jackson *et al.* (2008b) provides motivation and a basis for construction of improved physical models.

## 2. Method

To test the hypothesis that tides have been responsible for reducing  $e$ , we numerically integrated the canonical tidal evolution equations of Goldreich & Soter (1966) and Kaula (1968) backwards in time for all close-in planets for which we have adequate information (see Jackson *et al.* (2008a) for details). For each planet, we began the integration with the current best estimates of  $e$  and  $a$  ( $e_{current}$  and  $a_{current}$ ) and integrated backwards over the estimated age of the host star to find the orbital elements,  $e_{initial}$  and  $a_{initial}$ . We assumed tidal evolution dominated the orbital evolution after the protoplanetary disk dissipated and collisional effects became negligible, which probably happened only a few Myr after the star's formation. Hence  $e_{initial}$  and  $a_{initial}$  may describe the orbits at that time. We repeated the integration for 289 combinations of  $Q_p$  and  $Q_*$ , each  $Q$  ranging from  $10^4$  to  $10^8$ .

Our study involves a number of assumptions and approximations which are detailed and discussed by Jackson *et al.* (2008a). In particular, our model assumes host stars rotate much more slowly than their close-in companion planet revolves, an assumption largely corroborated by observation (Trilling (2000); Barnes (2001)). As a result, tides raised on the star by the planet tend to decrease both  $e$  and  $a$ , as do tides raised on the planet. Our results will inevitably need to be revisited and refined as improved data and physical models become available. Of necessity, we considered only planets for which the reported  $e$  is non-zero, and for which there is some estimate available for the age of the system. Even with these restrictions, we can still study about 40% of all known close-in planets. Stellar masses and radii come from a variety of sources: Bakos *et al.* (2007a); Bakos *et al.* (2007b); Da Silva *et al.* (2006); Fischer & Valenti (2005); Moutou *et al.* (2006); Saffe *et al.* (2006); Gorda & Svechnikov (1996); and Takeda *et al.* (2007). For planetary masses, we use the radial-velocity minimum mass. For planetary radii, if  $M_p > 0.3$  Jupiter's mass, we fix  $R_p = 1.2 R_J$  since, for Jovian planets in this range of mass, the

radius is insensitive to mass (Hubbard (1984)). This value is near the average radius of almost all observed transiting extra-solar planets. For planets with  $M_p < 0.3$  Jupiter's mass, we assume the planet has Jupiter's density and scale  $R_p$  accordingly. This assumption agrees fairly well with the observed radius for GJ 436 b (Gillon *et al.* (2007); Deming *et al.* (2007)). Where  $R_p$  and  $M_p$  have been determined directly from transit observations, we instead use those values. Planetary data were taken from a variety of sources: Bakos *et al.* (2007a); Bakos *et al.* (2007b); Butler *et al.* (2006); Da Silva *et al.* (2006); Deming *et al.* (2007); Gillon *et al.* (2007); Johnson *et al.* (2006); Knutson *et al.* (2007); Laughlin *et al.* (2005); Lovis *et al.* (2006); McArthur *et al.* (2004); Maness (2007); Mayor (2004); Moutou *et al.* (2006); Rivera *et al.* (2005); Udry *et al.* (2002); Valenti & Fischer (2005); Vogt *et al.* (2005); Wright *et al.* (2006); and Zucker *et al.* (2004).

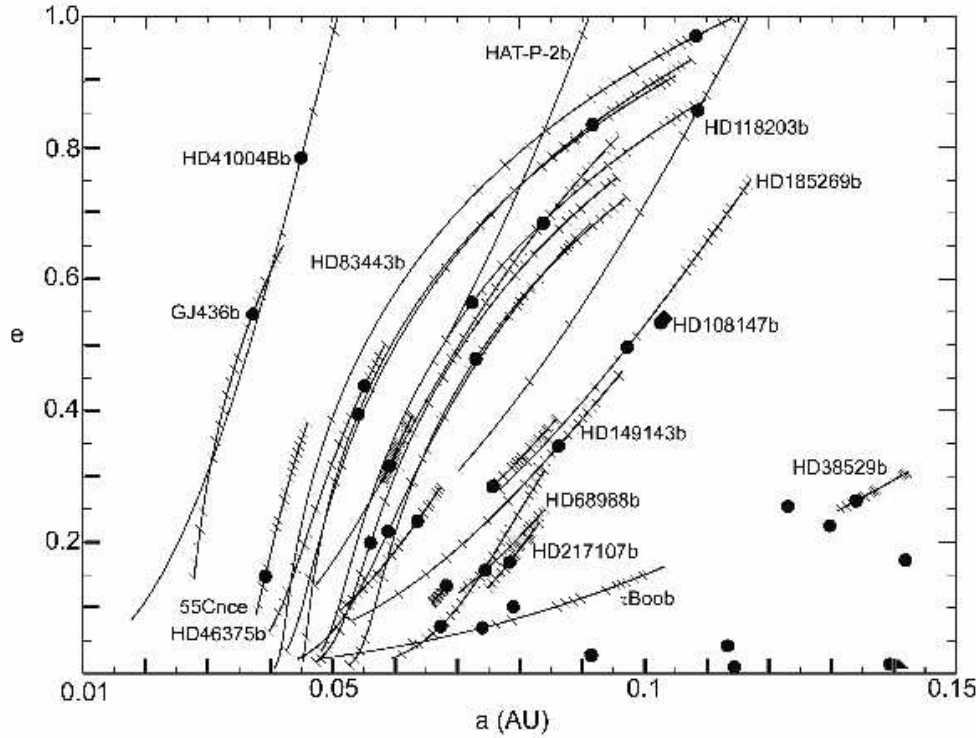
Based upon the above model, we can also calculate the tidal heating rate for planets undergoing tidal circularization. The tidal heating of the planet results from the reduction of orbital energy (and hence  $a$ ) due to energy dissipation in the planet. By tracking changes in  $a$  due to tides raised on the planet throughout the process of tidal circularization, we can estimate the tidal heating rate over the whole lifetime of the planet (Peale *et al.* (1979)). The magnitude and, indeed, the very shape of the past heating curve over time depends sensitively upon the assumed current orbital elements. Accordingly, we have calculated multiple plausible heating curves for each planet, corresponding to the range of observationally allowed values for  $e_{current}$  and  $a_{current}$ .

### 3. Orbital Evolution

The 289 combinations of  $Q_p$  and  $Q_*$  that we tested gave a wide variety of distributions of initial eccentricities for the close-in planets. (Remember “initial” refers to the time, shortly after formation, that a planet’s orbital evolution begins to be dominated by tides.) In addition to the current orbits (squares), Figure 1 shows the computed distributions of  $e_{initial}$  and  $a_{initial}$  (filled triangles) for the case of  $Q_p = 10^{6.5}$  and  $Q_* = 10^{5.5}$ . We compared the  $e_{initial}$  distributions for the close-in planets with the standard  $e$  distribution observed for  $a > 0.2$  AU. In this case, the agreement is excellent, with a K-S score of 90%. These  $Q$  values are well within plausible ranges. Reasonable fits can also be obtained with other values of  $Q_*$  as long as  $Q_p \sim 10^{6.5}$ . Other good fits (K-S  $\sim 70\%$ ) have  $Q_* = 10^{4.25}$  or  $> 10^{6.75}$ .

Figure 2 shows the evolutionary track of  $a$  and  $e$  over time for each of the planets, in the case of the best-fit  $Q$  values. The current orbital elements are at the lower left end of each trajectory in  $(a, e)$  space. These points correspond to empty squares in Figure 1. The tick marks show the orbital elements at intervals of 500 Myr, going back in time, for 15 Gyr, from the present toward the upper right. Black dots have been placed at a point representing the best age estimate for the planetary system. These same points appear as the triangles in Fig. 1.

The evolutionary histories derived here include substantial changes in semi-major axis coupled with the changes in eccentricity. For many close-in planets, Figures 1 and 2 show that initial  $a$  values were significantly higher than the currently observed values. These initial values of  $a$  likely represent their locations at the termination of gas disk migration in each early planetary system. Given the extent of tidal migration for observed close-in planets, planets that formed inward of 0.04 AU could have subsequently fallen into their host stars. Consideration of such hypothesized scenarios will be the subject of future work.



**Figure 2.** Tidal evolution of  $e$  and  $a$  for the sample of known close-in extra-solar planets using our best-fit values of  $Q_* = 10^{5.5}$  and  $Q_p = 10^{6.5}$ . Solid curves represent the trajectories of orbital evolution from current orbits (lower left end of each curve) backward in time (toward the upper right). On the trajectories, tick marks are spaced every 500 Myr to indicate the rate of tidal evolution. Tidal integrations were performed for 15 Gyr for all planets, but the filled circles indicate the initial values of orbital elements at the beginning of each planet's life. Due to space restrictions, most planets are not labeled, however they can be identified by the  $(e, a)$  values at the lower left end of each trajectory. Jackson *et al.* (2008a) includes a table of  $e_{\text{current}}$  and  $a_{\text{current}}$  (Table 1) which can be used to identify planets in this figure.

#### 4. Tidal Heating

Using the results of our tidal evolution calculations, we modeled the past tidal heating rates for numerous planets. In Figure 3, we illustrate plausible heating histories for three interesting examples: HD 209458 b, HAT-P-1 b and GJ 436 b, using  $Q_p = 10^{6.5}$  and  $Q_* = 10^{5.5}$ . Results for other planets are presented in Jackson *et al.* (2008b). Note that in these examples the planetary radii have been held constant, even though the changing heating rates would most likely produce a changing radius.

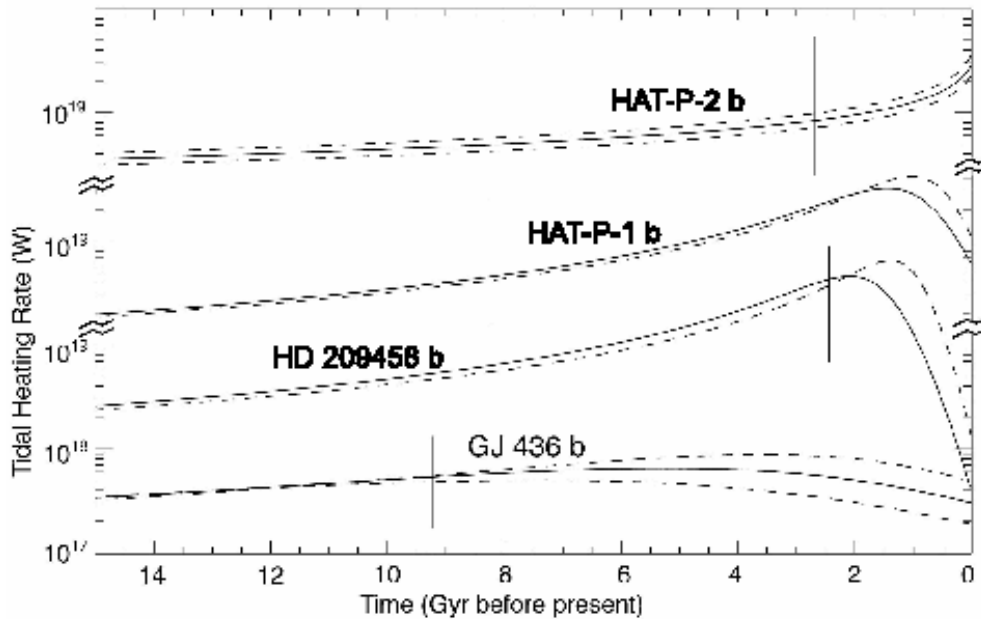
For HD 209458 b, Burrows *et al.* (2007) suggest that a heating rate of about  $4 \times 10^{19}$  W would be required to yield the observed planetary radius, which is much larger than any allowable current tidal heating rate. However, the history plotted in Figure 3 shows that the required heating rate was available as recently as 1 billion years ago. If the lag in the response of the planet to the heating rate were on the order of a billion years, this heating rate may explain the observed large radius. Such a lag seems reasonable based on the long duration of the influence of heat of formation on the planet's radius in the modeling by Burrows *et al.* (2007).

Like HD 209458 b, HAT-P-1 b's observed radius of  $1.36 R_J$  (Bakos *et al.* (2007a)) is larger than expected from theoretical modeling that did not include tidal heating

(Guillot (2005)). Similar to HD 209458 b, Figure 3 shows that its heating rate  $\sim 1$  Gyr ago was substantially higher than the present tidal heating. For both HD 209458 b and HAT-P-1 b, the substantial heating rate  $\sim 3\text{--}4 \times 10^{19}$  W about 1 Gyr ago may help account for the discrepancy between the large observed planetary radii and the predictions of physical modeling.

In the case of HAT-P-2 b there also has been a substantial amount of tidal heating. The current heating rate is similar to the maximum rate attained by HD 209458 b and HAT-P-1 b, so again we might expect a larger radius than predicted by theory that ignored tidal heating. In this case, however, the measured radius is actually about 12% smaller than predicted (Bakos *et al.* (2007b)). Thus there is a discrepancy between theory and observation even if tidal heating is neglected. The fact that there is likely a high rate of tidal dissipation makes the problem even worse. On the other hand, a key factor in the reconciliation may be that, while the current tidal heating rate is high and increasing, in the recent past the heating rate was much lower. HAT-P-2 b is still on the increasing part of the heating curve, which is unusual among planets considered here, most of which have passed their peaks. The fact that the heating rate was several times smaller a billion years ago than it is now may help explain the small radius.

Figure 3 also illustrates tidal heating histories for GJ 436 b. This planet has a measured



**Figure 3.** The tidal heating rates for planets HAT-P-2 b, HAT-P-1 b, HD 209458 b, and GJ 436 b as a function of time. The present time ( $t = 0$ ) is at the right. The solid vertical line through each of the heating curves represents the estimated time of formation for that planet. Note that the vertical scales have been shifted for HAT-P-2 b and HAT-P-1 b to make its curves more visible. (The vertical scale that corresponds to each curve is the scale intersected by that curve.) The solid curve for each planet is based on the current nominal eccentricity value. The dashed lines represent upper and lower limits on the heating, *i.e.* they assume the extreme values of current  $a$  and  $e$  consistent with observations that give the maximum and minimum current tidal heating rates, respectively. For HAT-P-1 b and HD 209458 b, observations could not exclude a current eccentricity of zero, so the lower bound on heating rates is formally zero. Hence in those cases only one dashed line is shown, representing the upper limit.

radius consistent with theoretical models, independent of tidal heating (Gillon *et al.* (2007)). The tidal heating history shown in Figure 3 is consistent with that result. Compared with the previous two cases, the maximum heating rate was two orders of magnitude less, small enough perhaps not to affect the radius.

## 5. Discussion

This investigation supports the hypothesis that tidal interactions between a star and a planet are responsible for the relatively small  $e$  values of close-in planets (Rasio *et al.* (1996)), although our calculations introduce important corrections to previous studies. Because even close-in planets evidently formed with much larger  $e$ , the processes that governed their early dynamics were probably similar to other planets. A plausible mechanism for producing the initial  $e$  distribution of extra-solar planets is planet-planet scattering (Rasio & Ford (1996); Weidenschilling & Marzari (1996)), although some modification to the original model are needed (Barnes & Greenberg (2007)).

If the tidal circularization hypothesis is borne out, then our study also provides constraints on  $Q$  values. We find that agreement between eccentricity distributions requires  $Q_p \sim 10^{6.5}$ , with  $Q_* \sim 10^{4.25}$ ,  $10^{5.5}$  (best fit), or  $10^{6.75}$ . This  $Q_p$  is in good agreement with other constraints (Yoder & Peale (1981); Ogilvie & Lin (2004)), and this  $Q_*$  agrees well with studies of binary star circularization (Mathieu (1994)). Of course, it is likely that individual planets and stars have unique  $Q$  values, owing to variation in their internal structures. These  $Q$  values are only meant to be representative values, good for the population as a whole. Corrections to our tidal model might also result in different suitable  $Q$  values.

Significant reductions in semi-major axes have accompanied the changes in eccentricity, with important implications. First, models of protoplanetary migration in the primordial gas disk need not carry “hot Jupiters” in as far as their current positions. Lin *et al.* (1996) proposed that migration in the gas disk halted near the inner edge of the disk, a boundary determined by clearing due to the host star’s magnetosphere. Our results show that the inner edge was probably farther out than indicated by the current semi-major axes of the planets, which were only reached during tidal migration long after the nebula had dissipated. In order to evaluate where migration due to the gas disk halted, (and thus where the inner edge of the nebula was) models should account for the subsequent tidal evolution.

The tidal changes in orbital semi-major axes also have implications for observations of planetary transits, such as surveys of young open galactic clusters (Bramich *et al.* (2005); von Braun *et al.* (2005); Burke *et al.* (2006)). The probability to observe a planetary transit increases for smaller semi-major axes, but decreases as orbits become more circular (Borucki & Summers (1984); Barnes (2007)). Tidal evolution means that the probability of an observable transit depends on a star’s age, but the exact relation depends on the particular evolutionary path through  $(a, e)$  space. As our understanding of the statistics of tidal evolution paths improves, the observed frequency of transits in the field and in open clusters may eventually help to constrain planetary formation scenarios, distinguishing, for example, between the relative roles of embedded migration and of gravitational scattering, which set up the initial conditions for tidal evolution. Transit statistics may not yet be refined enough to be sensitive to detect this effect (Pepper & Gaudi (2006)), but such systematic effects may show up in future surveys.

The tidal heating calculations here suggest that past tidal heating may well have played an important role in the evolution of the physical properties of many extra-solar planets, specifically the planetary radius. We caution that the specific calculations displayed here

depend on numerous assumptions and several uncertain parameters. The heating rates correspond to the orbital evolution trajectories computed by Jackson *et al.* (2008a), and various caveats are discussed in detail there. It is quite likely that the actual thermal history of any particular planet was different to some degree from what we show here. In particular, when analyzing radial-velocity observations to solve for a planet's orbit, it is difficult to rule out a completely circular orbit, in which case the tidal heating rate would have been zero. However, the unavoidable point is that past tidal heating may be significant for many planets and should be considered as a factor in theoretical modeling of physical properties of exoplanets.

For every planet whose tidal evolution we modeled, we have calculated corresponding tidal heating histories (Jackson *et al.* (2008b)). For the cases presented here, see Fig. 3, we see that past tidal heating may provide a previously unconsidered source of heat for planets with larger-than-predicted radii. However, it may make things worse in cases where measured radii seemed to fit the current models. Theoretical models of tidal evolution and planetary interiors will generally need to be adjusted and improved so as to yield a match between predicted and observed planetary radii.

To conclude, we find that the distribution of orbital eccentricities for exoplanets was once strikingly uniform across all semi-major axes. By varying tidal dissipation parameters, we can match the original distribution of close-in planetary eccentricities to that of planets far from their host star for stellar and planetary  $Q$ 's  $\sim 10^{5.5}$  and  $\sim 10^{6.5}$ , which are consistent with previous estimates. After the formation of the close-in exoplanets, tides raised on the host star and on the planet acted over Gyrs to reduce orbital eccentricities and semi-major axes of the close-in exoplanets. This reduction in  $e$  as well as  $a$  has important implications for the thermal histories of close-in exoplanets and for transit studies.

We also find that tidal heating in the past was significantly larger than current heating. For example, about 1 Gyr ago, HD 209458 b may have undergone tidal heating 100 times the present value. This substantial heating may help resolve the mystery of the anomalously large radii observed for many transiting planets today. If the lag in response of the planetary radii to tidal heating is of order a Gyr, then past tidal heating must be included in models of exoplanetary radii. Previous studies suggest such a lag is reasonable. However, further studies are required to elucidate this effect.

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